

# Beyond Two Dark Energy Parameters

Devdeep Sarkar

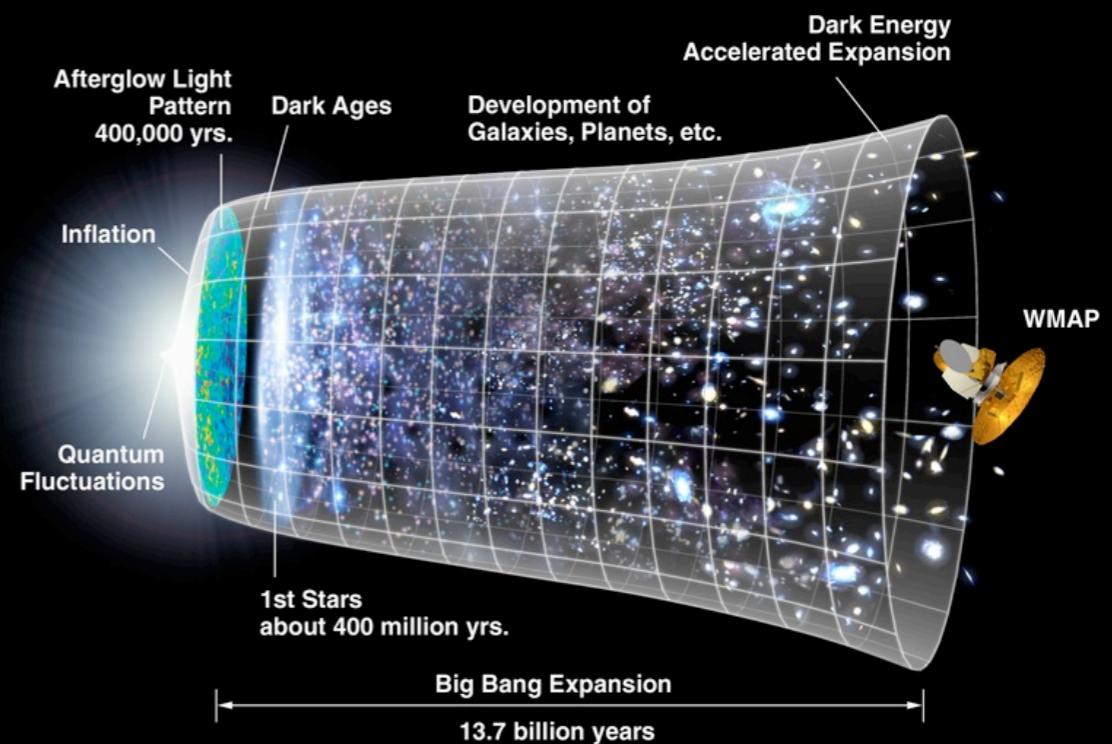
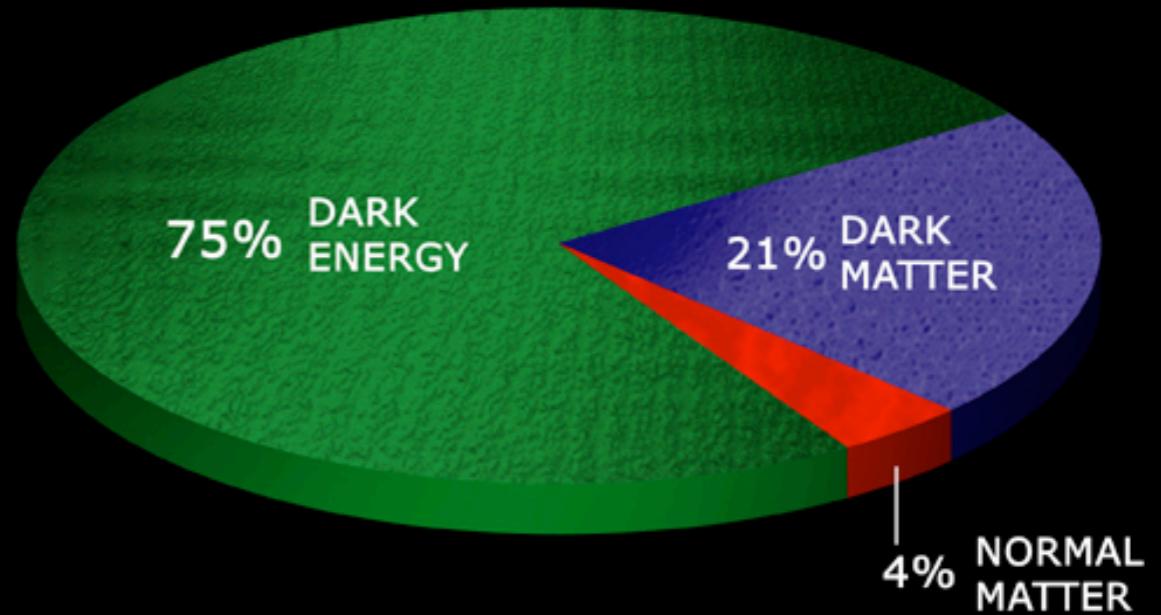
in collaboration with:

Scott Sullivan (UCI/UCLA), Shahab Joudaki (UCI), Alexandre Amblard (UCI),  
Daniel Holz (Chicago/LANL), and Asantha Cooray (UCI)

# Energy Budget of the Universe

$$\Omega_{Total} = \Omega_b + \Omega_{dm} + \Omega_{de}$$
$$\approx 1$$

$$\left( \Omega_X \equiv \frac{\rho_X}{\rho_{crit}} \right)$$



Puzzle! Puzzle! Puzzle!

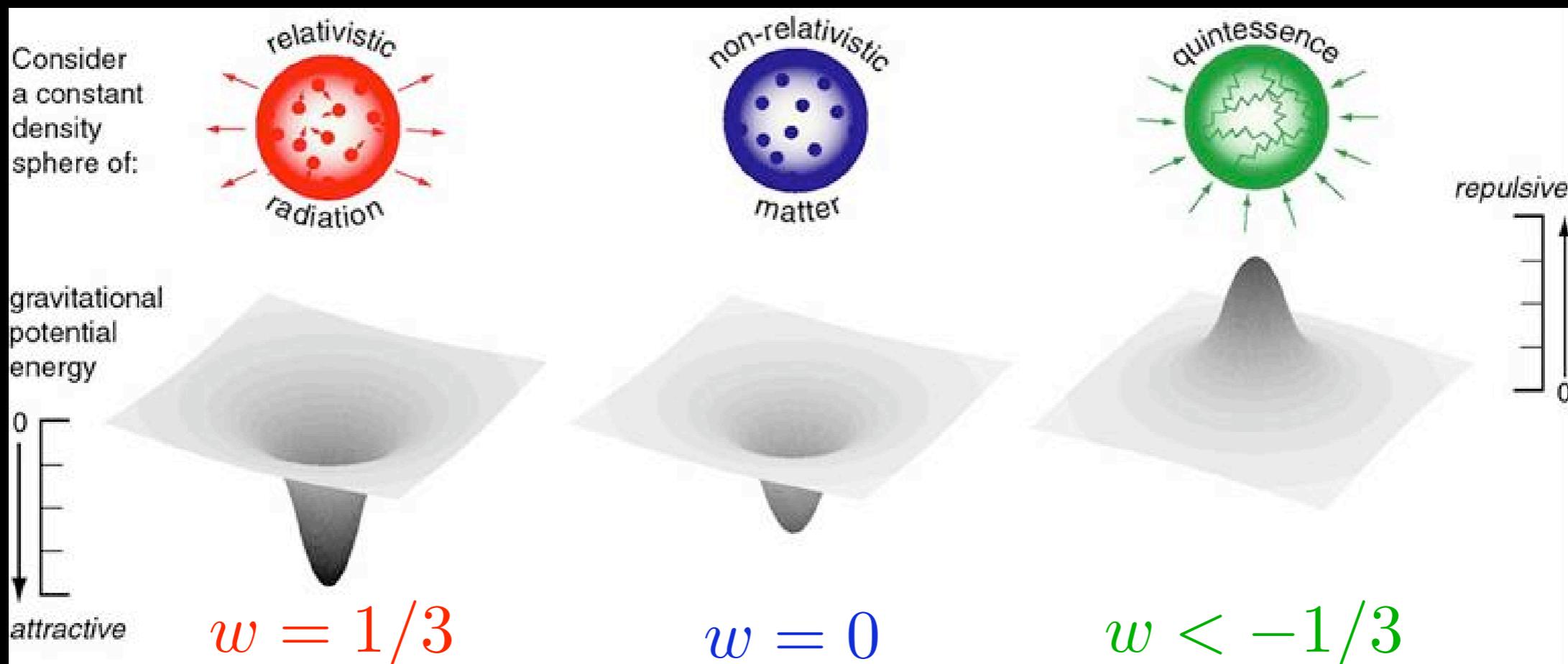
# What is Dark Energy?

“Dark Energy is made from an exclusive blend of vital L-amino acids, beneficial vitamins and bionutrients that allows faster and greater ion penetration of the cell walls, visibly enhancing the rate of growth”



# Dark Energy Equation Of State

$$T_{\mu}^{\nu} = \text{diag}(\rho, -p, -p, -p) \quad p = w\rho$$



For Cosmological Constant...  $w = -1$

# “Seeing” The Dark Energy

# “Seeing” The Dark Energy

...via its effect on the expansion of the Universe

$$H(z) = H_0 \left[ \Omega_m (1+z)^3 + \Omega_k (1+z)^2 + (1 - \Omega_k - \Omega_m) F(z) \right]^{1/2}$$

# “Seeing” The Dark Energy

...via its effect on the expansion of the Universe

$$H(z) = H_0 \left[ \Omega_m (1+z)^3 + \Omega_k (1+z)^2 + (1 - \Omega_k - \Omega_m) F(z) \right]^{1/2}$$

Approaches...

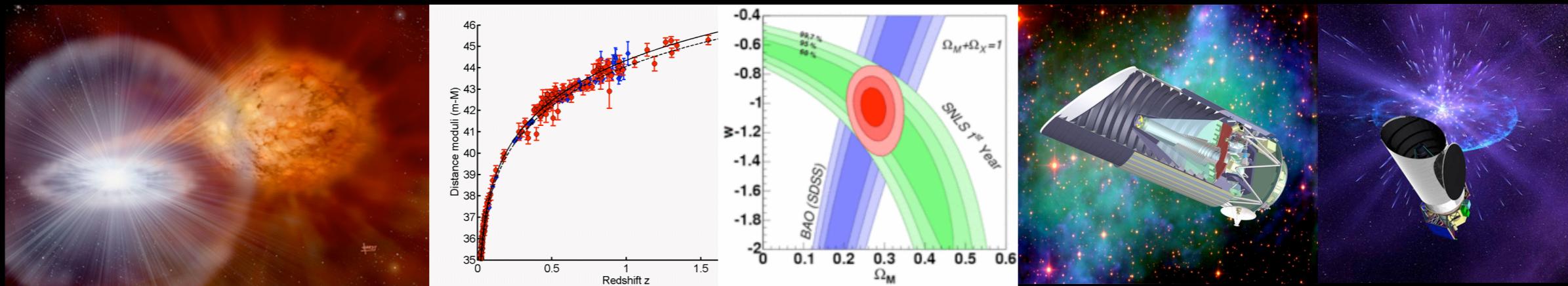
# “Seeing” The Dark Energy

...via its effect on the expansion of the Universe

$$H(z) = H_0 \left[ \Omega_m (1+z)^3 + \Omega_k (1+z)^2 + (1 - \Omega_k - \Omega_m) F(z) \right]^{1/2}$$

Approaches...

(I) Standard Candles: Luminosity Distance of SNe



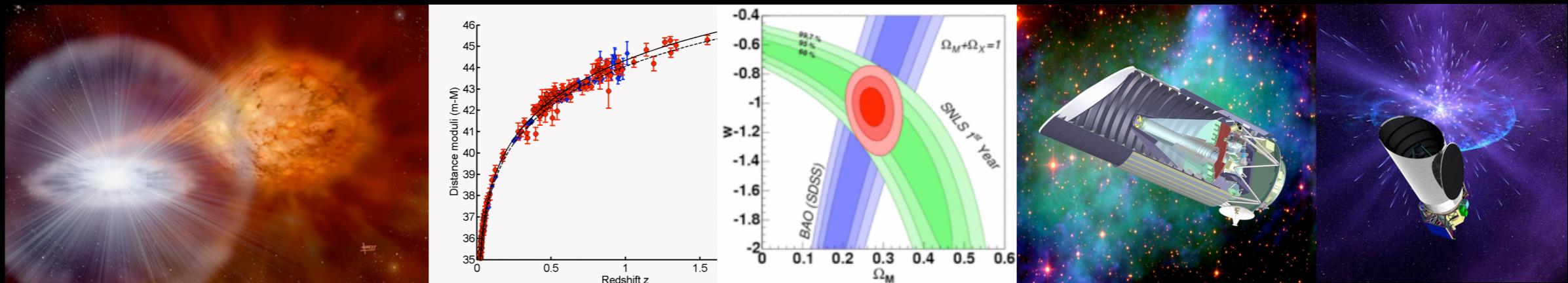
# “Seeing” The Dark Energy

...via its effect on the expansion of the Universe

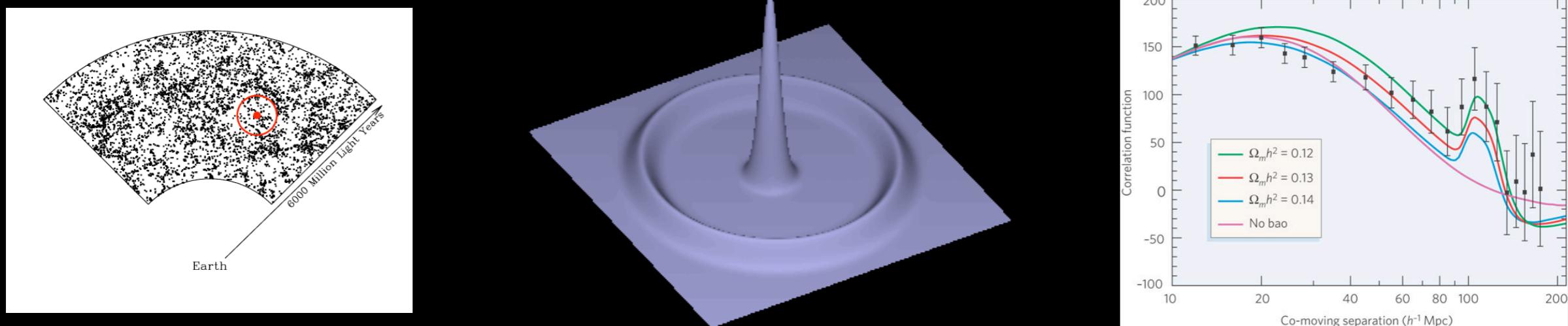
$$H(z) = H_0 \left[ \Omega_m (1+z)^3 + \Omega_k (1+z)^2 + (1 - \Omega_k - \Omega_m) F(z) \right]^{1/2}$$

Approaches...

(1) Standard Candles: Luminosity Distance of SNe



(2) Standard Rulers: Angular Diameter Distance via BAO



# DE EOS Revisited: Different Approaches...

# DE EOS Revisited: Different Approaches...

(A) Parameterize  $w(z)$

[Adopted by the DETF]

$$w(a) = w_0 + (1 - a)w_a \quad (\text{Linder 2003})$$

# DE EOS Revisited: Different Approaches...

(A) Parameterize  $w(z)$

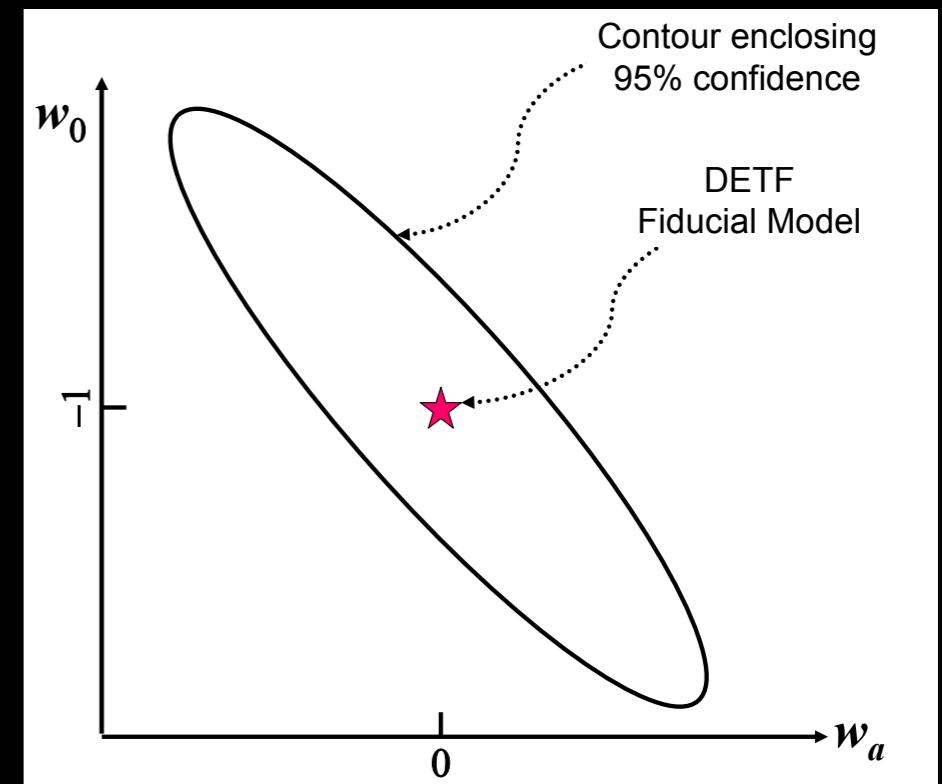
[Adopted by the DETF]

$$w(a) = w_0 + (1 - a)w_a$$

(Linder 2003)

**DETF Figure of Merit (FoM)**

Reciprocal of the Area of  
the Error Ellipse enclosing  
the 95% Confidence Limit  
in the  $w_0 - w_a$  Plane.



Larger Value of the FoM



Greater Accuracy

# DE EOS Revisited: Beyond Two Param...

## (B) Model-Independent Redshift-Binning of $w(z)$

# DE EOS Revisited: Beyond Two Param...

## (B) Model-Independent Redshift-Binning of $w(z)$

bin	$w_1(z)$	$w_2(z)$	$w_3(z)$	$w_4(z)$	$w_5(z)$	$w_6(z)$
z-ran	0-0.07	0.07-0.15	0.15-0.3	0.3-0.6	0.6-1.2	1.2-2.0

# DE EOS Revisited: Beyond Two Param...

## (B) Model-Independent Redshift-Binning of $w(z)$

$$F(z_n > z > z_{n-1}) = (1+z)^{3(1+w_n)} \prod_{i=0}^{n-1} (1+z_i)^{3(w_i - w_{i+1})}$$

# DE EOS Revisited: Beyond Two Param...

## (B) Model-Independent Redshift-Binning of $w(z)$

bin	$w_1(z)$	$w_2(z)$	$w_3(z)$	$w_4(z)$	$w_5(z)$	$w_6(z)$
z-ran	0-0.07	0.07-0.15	0.15-0.3	0.3-0.6	0.6-1.2	1.2-2.0

# DE EOS Revisited: Beyond Two Param...

## (B) Model-Independent Redshift-Binning of $w(z)$

bin	$w_1(z)$	$w_2(z)$	$w_3(z)$	$w_4(z)$	$w_5(z)$	$w_6(z)$
z-ran	0-0.07	0.07-0.15	0.15-0.3	0.3-0.6	0.6-1.2	1.2-2.0

Covariance Matrix

$$C = \langle \mathbf{w} \mathbf{w}^T \rangle - \langle \mathbf{w} \rangle \langle \mathbf{w}^T \rangle$$

# DE EOS Revisited: Beyond Two Param...

## (B) Model-Independent Redshift-Binning of $w(z)$

bin	$w_1(z)$	$w_2(z)$	$w_3(z)$	$w_4(z)$	$w_5(z)$	$w_6(z)$
z-ran	0-0.07	0.07-0.15	0.15-0.3	0.3-0.6	0.6-1.2	1.2-2.0

## (C) Decorrelated Estimates of $w(z)$

*(Huterer & Cooray 2005)*

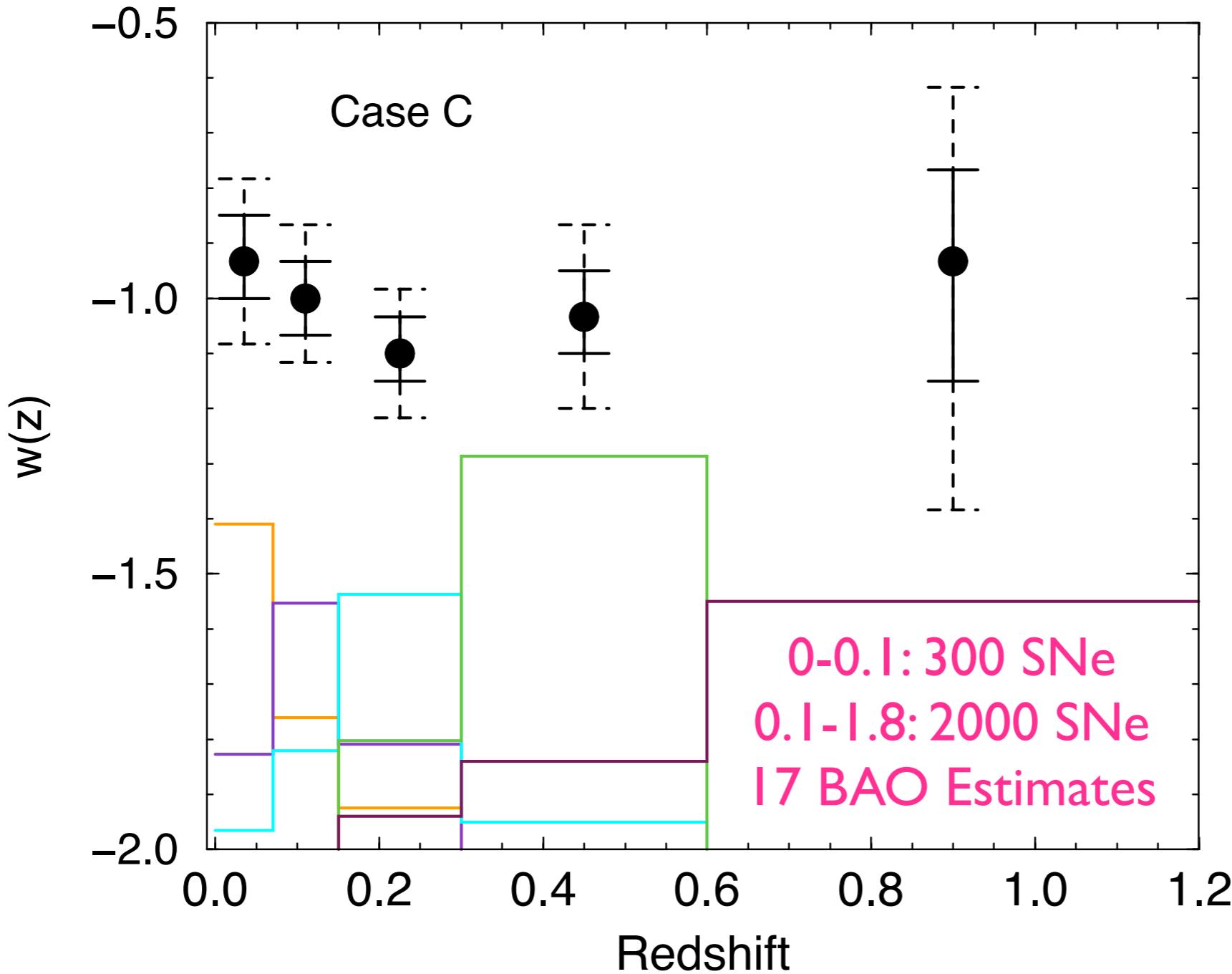
Diagonalize The Fisher Matrix

$$\mathbf{F} \equiv \mathbf{C}^{-1} = \mathbf{O}^T \boldsymbol{\Lambda} \mathbf{O}$$

# Our Analyses: Six Mock Scenarios

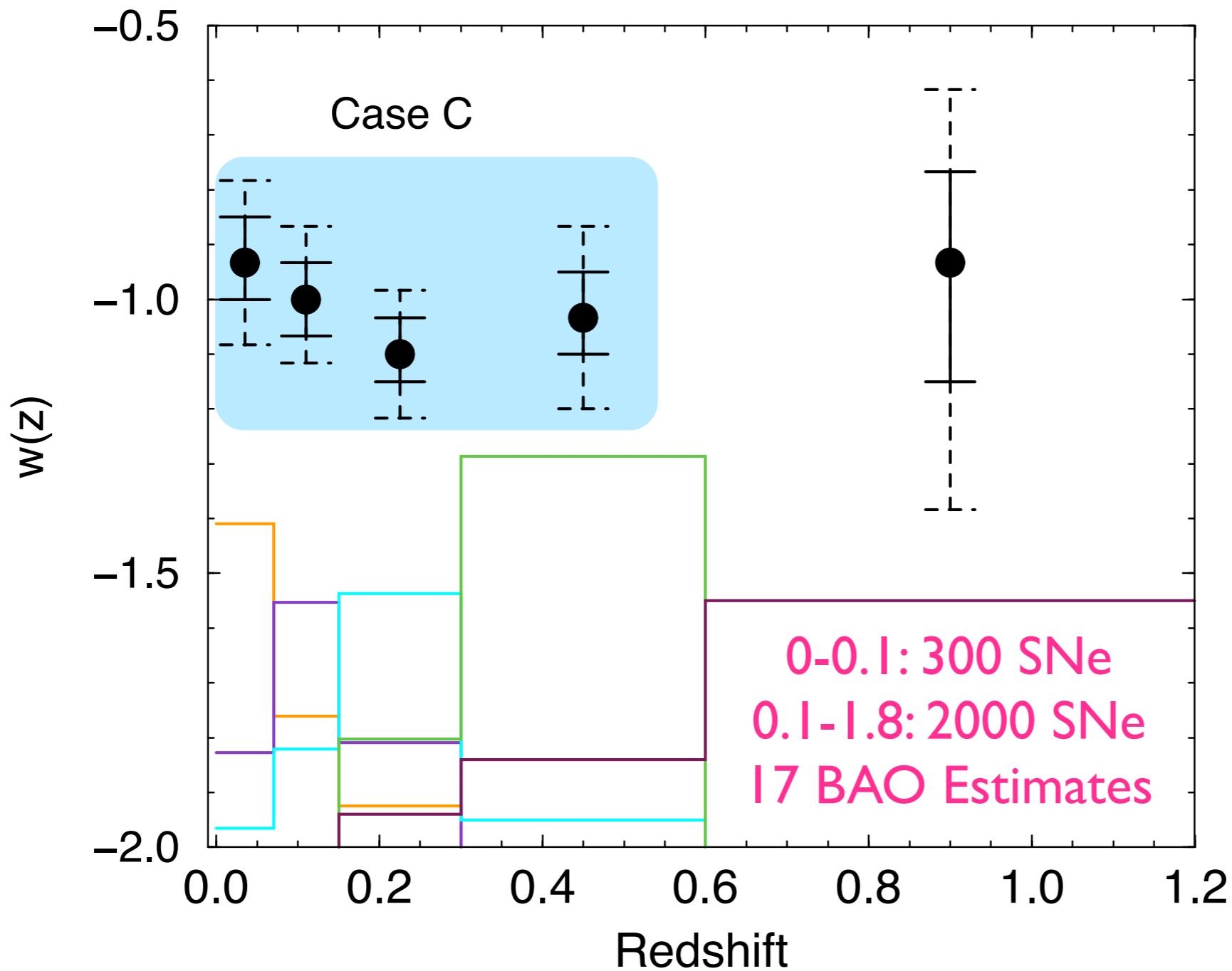
- Case A: 200 SNe up to  $z=1.8$ ; 2 BAO estimates
- Case B: 300 SNe out to  $z=0.1$  & 2000 SNe in the range  $0.1 < z < 1.8$ ; 7 BAO estimates
- Case C: Same as (B)... with 10 new BAO constraints
- Case D: 10,000 SNe out to  $z=2.0$ , 7 BAO estimates
- Case E: 10,000 SNe out to  $z=2.0$ ; 17 BAO estimates
- Case F: 200 SNe as in (A); 17 BAO estimates

# Results



$w_1$	0-.07
$w_2$	.07-.15
$w_3$	0.15-.3
$w_4$	0.3-0.6
$w_5$	0.6-1.2
$w_6$	1.2-2.0

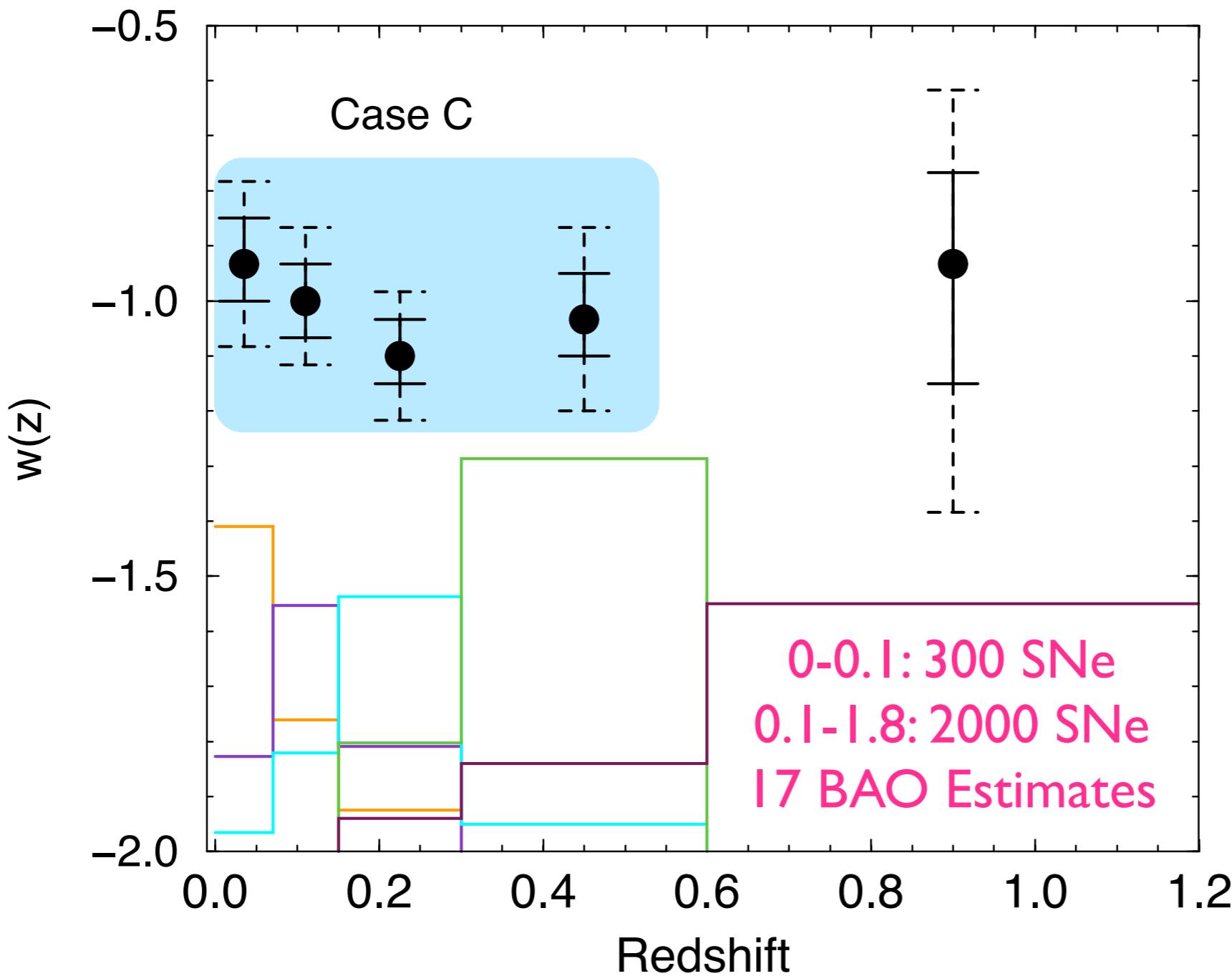
# Results



$w_1$	0-.07
$w_2$	.07-.15
$w_3$	0.15-.3
$w_4$	0.3-0.6
$w_5$	0.6-1.2
$w_6$	1.2-2.0

# Results

$$\text{FoM} = \left[ \sum_i \frac{1}{\sigma^2(w(z_i))} \right]^{1/2}$$



# Conclusion

- Future surveys can be expected to estimate three or more independent dark energy EOS parameters to an accuracy better than 10%
- We propose a parameterization-independent Figure of Merit (FoM):

$$\text{FoM} = \left[ \sum_i \frac{1}{\sigma^2(w(z_i))} \right]^{1/2}$$

- The FoM holds for an arbitrary # of dof

.com

THANK YOU



# Extra Slides

# Constraining with Standard Candles (SNe)

$$\begin{aligned}m - M &= -2.5 \log_{10} \left[ \frac{\mathcal{L}/4\pi d_L^2}{\mathcal{L}/4\pi(10\text{pc})^2} \right] \\&= 5 \log_{10} \left( \frac{d_L}{\text{Mc}} \right) + 25\end{aligned}$$



# Constraining with Standard Candles (SNe)

$$\begin{aligned}m - M &= -2.5 \log_{10} \left[ \frac{\mathcal{L}/4\pi d_L^2}{\mathcal{L}/4\pi(10\text{pc})^2} \right] \\&= 5 \log_{10} \left( \frac{d_L}{\text{Mc}} \right) + 25\end{aligned}$$

where



# Constraining with Standard Candles (SNe)

$$\begin{aligned}m - M &= -2.5 \log_{10} \left[ \frac{\mathcal{L}/4\pi d_L^2}{\mathcal{L}/4\pi(10\text{pc})^2} \right] \\&= 5 \log_{10} \left( \frac{d_L}{\text{Mc}} \right) + 25\end{aligned}$$

where



$$d_L(z) = (1+z) \frac{c}{H_0} \times \begin{cases} \frac{1}{\sqrt{|\Omega_k|}} \sinh \left( \sqrt{|\Omega_k|} \int_0^z \frac{dz'}{H(z')/H_0} \right) & \text{if } \Omega_k > 0, \\ \int_0^z \frac{dz'}{H(z')/H_0} & \text{if } \Omega_k = 0, \\ \frac{1}{\sqrt{|\Omega_k|}} \sin \left( \sqrt{|\Omega_k|} \int_0^z \frac{dz'}{H(z')/H_0} \right) & \text{if } \Omega_k < 0, \end{cases}$$

# Constraining with Standard Candles (SNe)

$$\begin{aligned}m - M &= -2.5 \log_{10} \left[ \frac{\mathcal{L}/4\pi d_L^2}{\mathcal{L}/4\pi(10\text{pc})^2} \right] \\&= 5 \log_{10} \left( \frac{d_L}{\text{Mc}} \right) + 25\end{aligned}$$

where

$$d_L(z) = (1+z) \frac{c}{H_0} \times \begin{cases} \frac{1}{\sqrt{|\Omega_k|}} \sinh \left( \sqrt{|\Omega_k|} \int_0^z \frac{dz'}{H(z')/H_0} \right) & \text{if } \Omega_k > 0, \\ \int_0^z \frac{dz'}{H(z')/H_0} & \text{if } \Omega_k = 0, \\ \frac{1}{\sqrt{|\Omega_k|}} \sin \left( \sqrt{|\Omega_k|} \int_0^z \frac{dz'}{H(z')/H_0} \right) & \text{if } \Omega_k < 0, \end{cases}$$

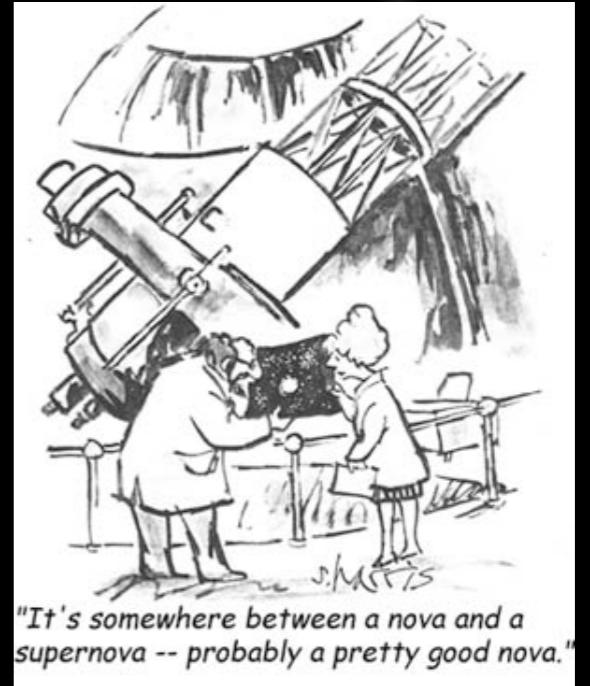
where



# Constraining with Standard Candles (SNe)

$$\begin{aligned}m - M &= -2.5 \log_{10} \left[ \frac{\mathcal{L}/4\pi d_L^2}{\mathcal{L}/4\pi(10\text{pc})^2} \right] \\&= 5 \log_{10} \left( \frac{d_L}{\text{Mc}} \right) + 25\end{aligned}$$

where



$$d_L(z) = (1+z) \frac{c}{H_0} \times \begin{cases} \frac{1}{\sqrt{|\Omega_k|}} \sinh \left( \sqrt{|\Omega_k|} \int_0^z \frac{dz'}{H(z')/H_0} \right) & \text{if } \Omega_k > 0, \\ \int_0^z \frac{dz'}{H(z')/H_0} & \text{if } \Omega_k = 0, \\ \frac{1}{\sqrt{|\Omega_k|}} \sin \left( \sqrt{|\Omega_k|} \int_0^z \frac{dz'}{H(z')/H_0} \right) & \text{if } \Omega_k < 0, \end{cases}$$

where

$$H(z) = H_0 \left[ \Omega_m (1+z)^3 + \Omega_k (1+z)^2 + (1-\Omega_k-\Omega_m) F(z) \right]^{1/2}$$

# DE EOS Revisited: Different Approaches...

(A) Parameterize  $w(z)$

[Adopted by the DETF]

$$w(z) = w_0 + (1 - a)w_a$$

(Linder 2003)

# DE EOS Revisited: Different Approaches...

(A) Parameterize  $w(z)$

[Adopted by the DETF]

$$w(z) = w_0 + (1 - a)w_a \quad (\text{Linder 2003})$$

And hence

# DE EOS Revisited: Different Approaches...

(A) Parameterize  $w(z)$

[Adopted by the DETF]

$$w(z) = w_0 + (1 - a)w_a \quad (\text{Linder 2003})$$

And hence

$$H(z) = H_0 \left[ \Omega_m(1+z)^3 + \Omega_k(1+z)^2 + (1 - \Omega_k - \Omega_m)F(z) \right]^{1/2}$$

# DE EOS Revisited: Different Approaches...

(A) Parameterize  $w(z)$

[Adopted by the DETF]

$$w(z) = w_0 + (1 - a)w_a \quad (\text{Linder 2003})$$

And hence

$$H(z) = H_0 \left[ \Omega_m(1+z)^3 + \Omega_k(1+z)^2 + (1 - \Omega_k - \Omega_m)F(z) \right]^{1/2}$$

where

# DE EOS Revisited: Different Approaches...

(A) Parameterize  $w(z)$

[Adopted by the DETF]

$$w(z) = w_0 + (1 - a)w_a \quad (\text{Linder 2003})$$

And hence

$$H(z) = H_0 \left[ \Omega_m(1+z)^3 + \Omega_k(1+z)^2 + (1 - \Omega_k - \Omega_m)F(z) \right]^{1/2}$$

where

$$F(z) = (1+z)^{3(1+w_0+w_a)} e^{-\frac{3w_a z}{(1+z)}}$$

# DE EOS Revisited: Different Approaches...

(A) Parameterize  $w(z)$

[Adopted by the DETF]

$$w(z) = w_0 + (1 - a)w_a$$

(Linder 2003)

And hence

$$H(z) = H_0 \left[ \Omega_m(1+z)^3 + \Omega_k(1+z)^2 + (1 - \Omega_k - \Omega_m)F(z) \right]^{1/2}$$

where

$$F(z) = (1+z)^{3(1+w_0+w_a)} e^{-\frac{3w_a z}{(1+z)}}$$

$$w_0 = -1 \text{ and } w_1 = 0$$



Cosmological  
Constant

# Different Scales

Redshift, z	Time, t (Years)	cz (km/s)	Distance (Mpc/h)	Temperature (K)
0	13.7 billion	0	0	2.725
0.01	13.5 billion	3000	30	2.752
0.033	13.1 billion	10,000	100	2.815
0.1	11.9 billion	30,000	300	2.998
1.0	04.8 billion	300,000	3000	5.450
4.0	~1 billion	1.2 million	12,000	13.625
100	400,000	0.3 billion	3.3 million	3000

$$1 \text{ Mpc} = 10^6 \text{ pc} = 3.26 \times 10^6 \text{ ly} = 3.08 \times 10^{22} \text{ m}$$